Numerical simulations of radiation processes in AGN.

Effect of chemical composition

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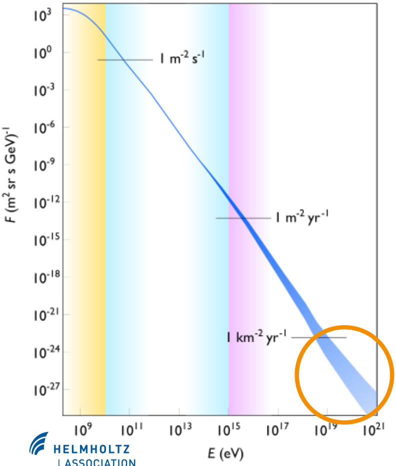
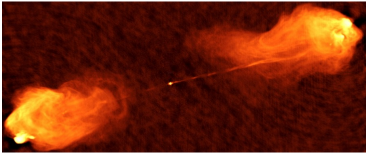
Introduction

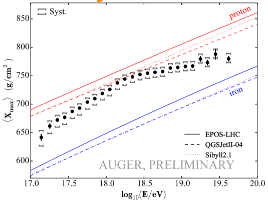
- **Question:** origin & composition of Ultra-High Energy Cosmic Rays
- Likely dominated by extra-galactic sources [Allard et al. 2007 astro-ph0512345]
- **Candidate:** Active Galactic Nuclei (AGN)





- Numerical model for hadronic interactions
- Predict cosmic-ray (CR) and ν spectra
- **Study effect of heavy elements**

motivation:
experimental data

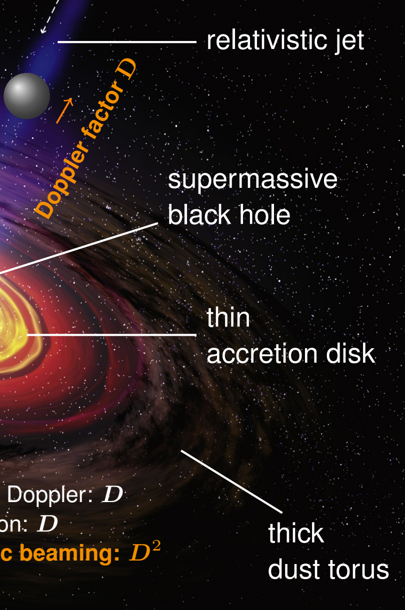


[Auger Collaboration, ICRC 2015]

Physical model

- CR acceleration in the jet
- Consider a plasma blob along the jet ⟶

  of radius $r = c\, t'_{\text{flare}}$
- $\boldsymbol{B}$ ⟶ **confinement of charged particles**
- $\mathbf{n}, \nu$ **escape directly**



- Useful transforms:

[Obs frame] [Jet rest frame]

Energy: $$E^{\text{obs}} = \left(\frac{D}{1+z}\right) E',$$

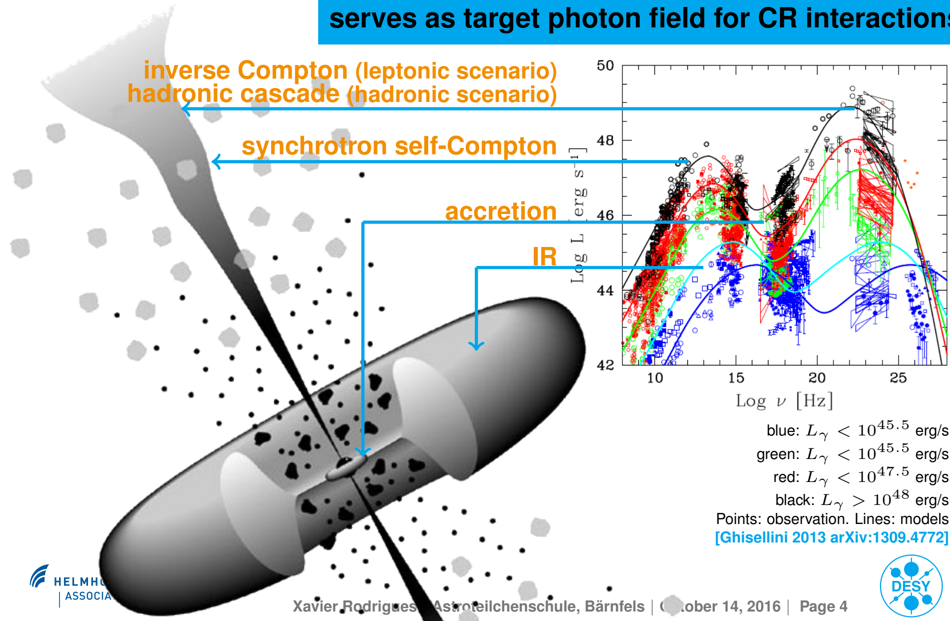
Luminosity: $$L^{\text{obs}} = \frac{D^4}{(1+z)^2} L'$$

relativistic Doppler: D
time dilation: D
relativistic beaming: D^2

Picture courtesy NASA, Dana Berry/Skyworks Digital

Blazar Spectral Energy Distribution (SED)

serves as target photon field for CR interactions



blue: $L_\gamma < 10^{45.5}$ erg/s

green: $L_\gamma < 10^{45.5}$ erg/s

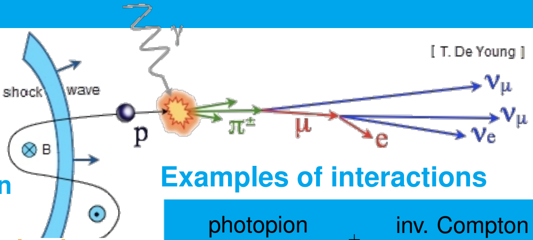
red: $L_\gamma < 10^{47.5}$ erg/s

black: $L_\gamma > 10^{48}$ erg/s

Points: observation. Lines: models

[Ghisellini 2013 arXiv:1309.4772]

Physical model – hadronic processes



Fermi-shock CR acceleration

↓

power-law CR injection

$$\frac{dn}{dE} = E^{-\alpha}$$

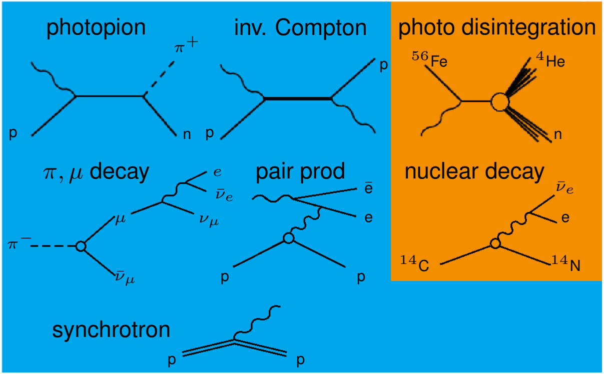
index α depends on shock compression ratio

(typically $\alpha = 2$)

Acceleration, synchrotron, escape time depend on magnetic field strength

Examples of interactions

If heavier nuclei are present,

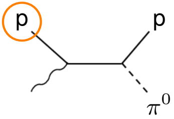


Interaction framework

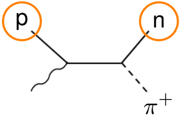
- **NeuCosmA** [Bearwald et al, AP 35 (2012) 508-529] numerical framework for nuclear cascades
- Developed for Gamma-Ray Bursts (GRB) ⟶ **apply to AGN**
- Numerically solve PDE system (1 eq. per particle species)

$$\frac{\partial N(E)}{\partial t} = \frac{\partial}{\partial E}\left(-b(E)\,N(E)\right) - \frac{N(E)}{t_{\mathrm{esc}}} + Q(E)$$

- $\frac{\partial N(E)}{\partial t}$ ↓ ==0 if steady state reached
- $\frac{\partial}{\partial E}\left(-b(E)\,N(E)\right)$ ↓ energy loss



- $\frac{N(E)}{t_{\mathrm{esc}}}$ ↓ particle escape
- $Q(E)$ ↓ particle injection



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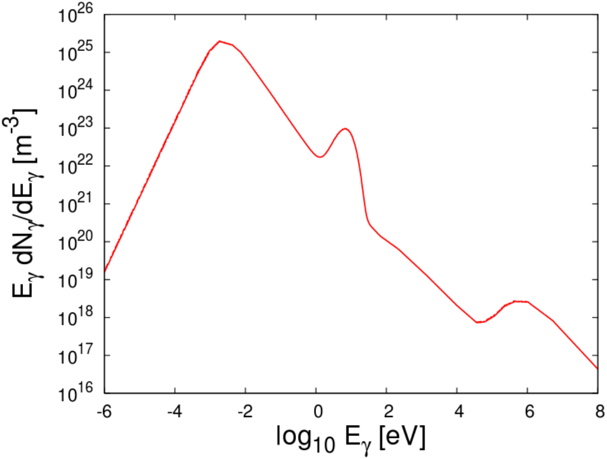
Preliminary results

- Input Blazar SED proposed in [Anchordoqui et al., AP 29 (2008) 1-13]



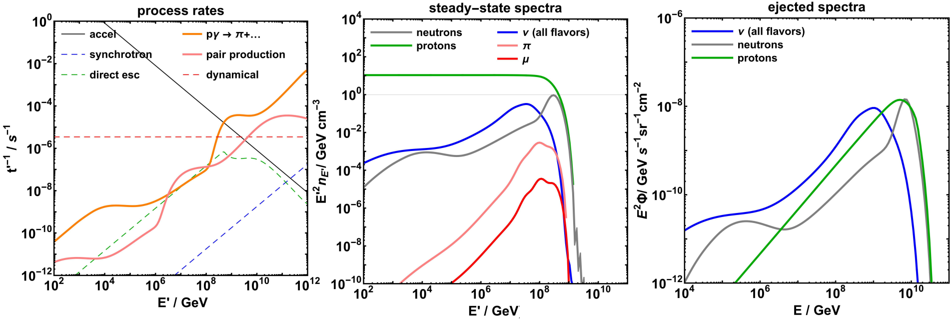
- $L'_\gamma = 10^{45}$erg/s
- $\Gamma = 10^{1.5}$
- $t^{obs}_{\text{flare}} = 10^4$ s

We will also consider

- $z = 0.1$
- $B = 0.9$ G

Preliminary results

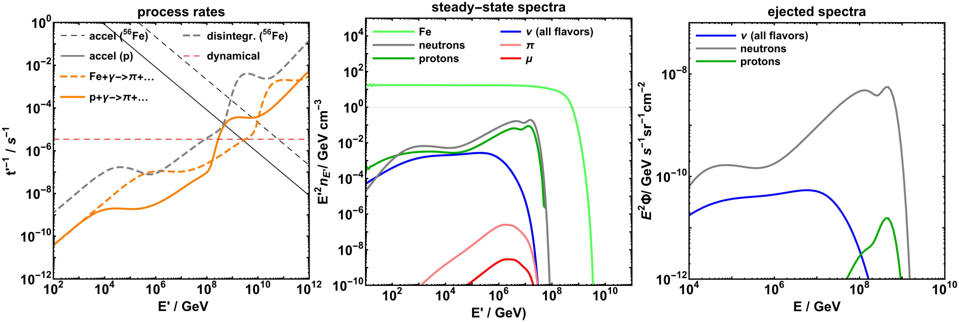
> Pure proton injection:



[X Rodrigues, A Fedynitch, D Boncioli, W Winter – in preparation]

Preliminary results

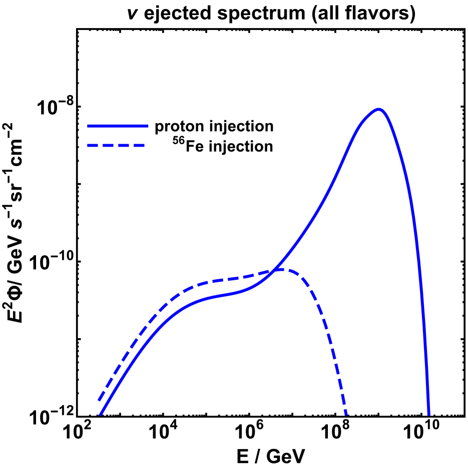
> **Pure iron-56 injection:**



[X Rodrigues, A Fedynitch, D Boncioli, W Winter – in preparation]

Preliminary results

> Comparison of total neutrino spectra

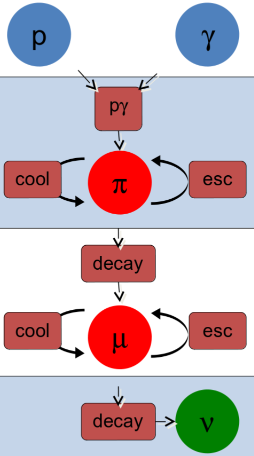


[X Rodrigues, A Fedynitch, D Boncioli, W Winter – in preparation]

Outlook

- Acceleration of heavy nuclei in AGN has a significant effect on the expected ν spectrum
- Currently attempting to reproduce other published results
- **Next step:** probe parameter space, particularly isotopic composition
- **In the future:** consider more sophisticated model to include photon feedback (self-consistent picture of photons–protons–nuclei)

Example: treatment of $p\,\gamma \;\rightarrow\; \pi + \cdots$ in NeuCosmA



Static species
or explicitly time-dependent
-- external input

X

Each variable species corresponds to one differential equation of the form:

Y

$$\frac{\partial N(E)}{\partial t} = \frac{\partial}{\partial E}\left(-b(E)\,N(E)\right) - \frac{N(E)}{t_{\mathrm{esc}}} + Q(E)$$

"Manual" species with injection only
Updated only on request
(e.g. in time-steps, where the data are written into a file)

Z

[W Winter, private note]